The energy density per unit volume of the atmosphere is the sum of internal, potential, and kinetic energies: $\rho(C_vT + gz + v^2/2)$.

Internal and potential energies are closely coupled. When air in a hydrostatic atmosphere is heated, it expands, doing work against the local pressure, and converting some of the heat into potential energy. Raising the temperature of a parcel by δT requires $C_v \delta T$ of internal energy and $R\delta T$ of work. The sum $(C_v + R)\delta T \equiv C_p \delta T$ is called the **enthalpy**.

The quantity which is conserved by a parcel moving in a hydrostatic atmosphere is the **dry static energy**. It includes the enthalpy rather than the internal energy. The energy flux is the product of the dry static energy and the velocity vector: $\rho(C_pT + gz + \frac{1}{2}v^2)\vec{v}$.

The conservation law for atmospheric energy density says that the rate of change of energy in a volume plus the flux out of the volume equals the diabatic heating rate Q

$$\frac{\partial}{\partial t} \{\rho(C_v T + gz + \frac{1}{2}v^2)\} + \nabla \cdot \{\rho(C_p T + gz + \frac{1}{2}v^2)\vec{v}\} = \rho Q$$

L. Li (LMD/CNRS):

Energy budget of the Earth and conversion of heat in different forms. The greenhouse effect of the atmosphere is due to its particular radiative properties: almost transparent for solar radiation, but almost opaque for the infrared terrestre radiation.



Radiative budget at the top of the atmosphere



Zonally and annually averaged of $\boldsymbol{v}\boldsymbol{T}$





Zonally and annually averaged of vZ





Zonally and annually averaged of vq





Zonally and annually averaged of vE





Northward VE transp. (stati) (J/kg m/s)









